

CRPropa Workshop 25.-27. September 2023

Ruhr University Bochum
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How is CR Propagation modelled?

Not within a single framework, because...

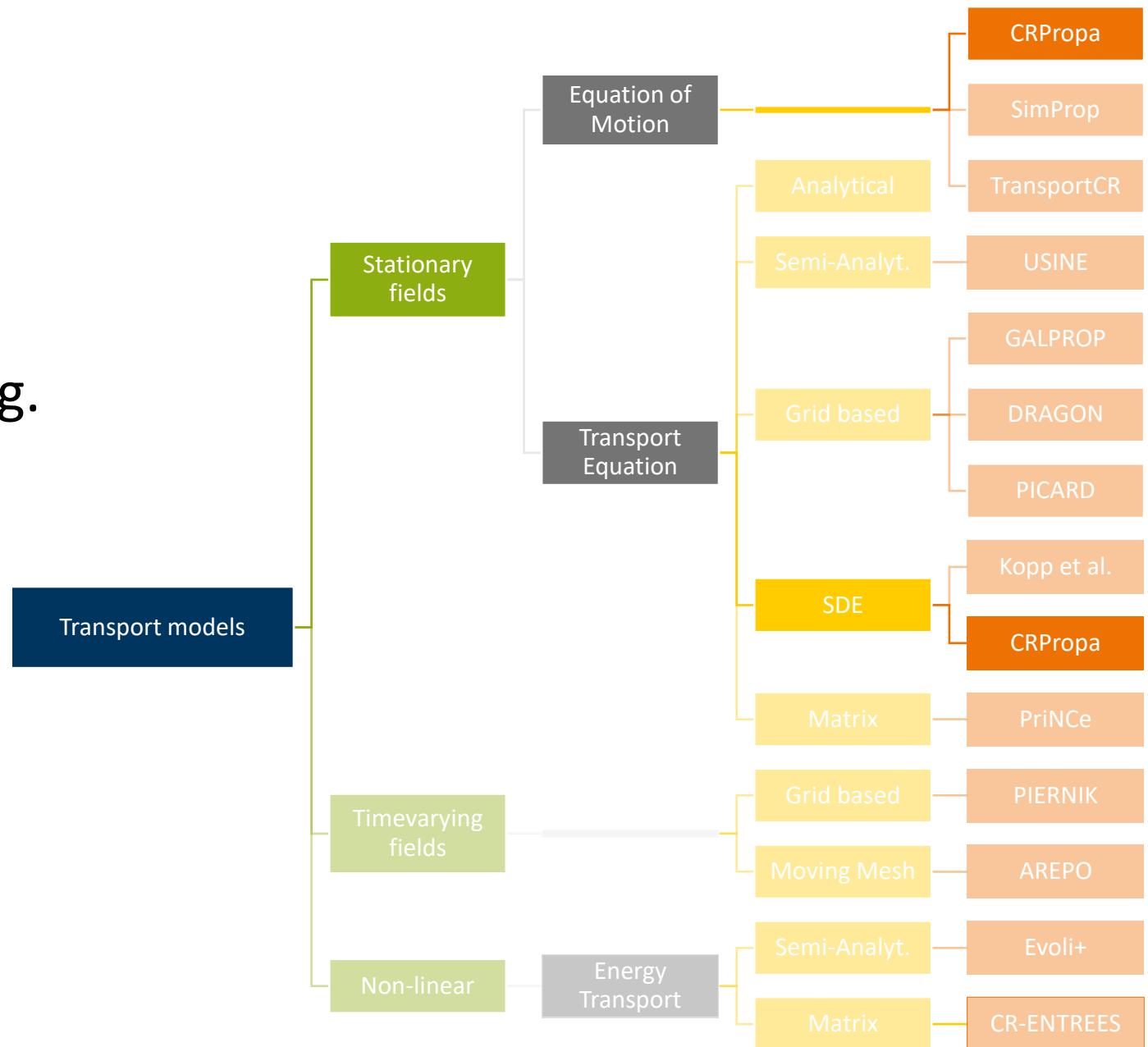
... the problem is too complicated.

Important questions:

- Are observables spatially resolved?
- Is time evolution important?
- Dominated by leptonic processes?
- Do cosmic rays influence their environment?
- What kind of targets are expected?
- Do the targets evolve in space or time?
- Coherent and/or turbulent magnetic fields?
- Is the particle transport ballistic or diffusive?

A family tree of propagation tools

...and many more are missing.



CRPropa

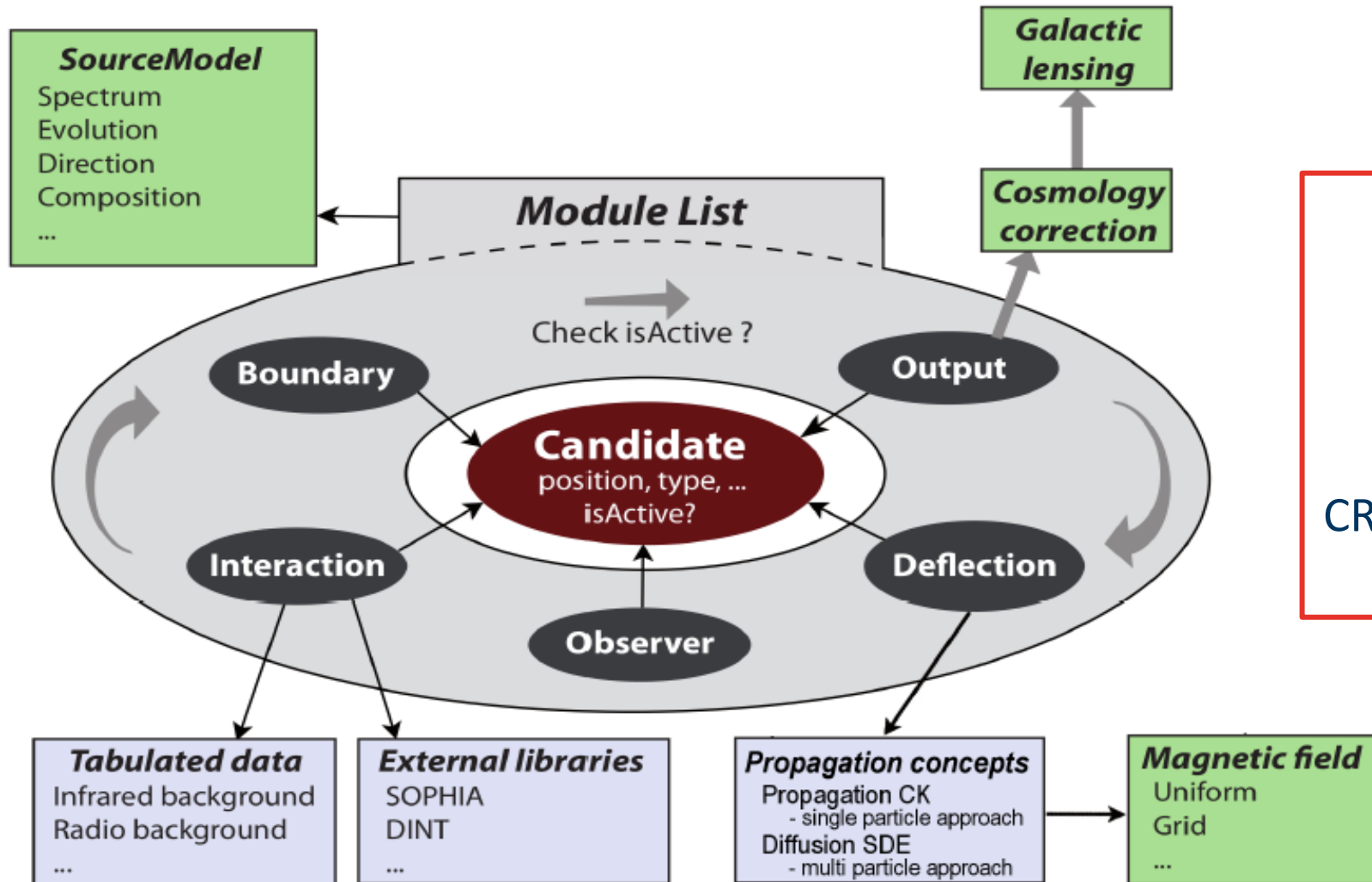
CRPropa – Cosmic Ray Propagation Framework

- Originally designed for 3D propagation of UHECRs in extra-galactic space
- Extended to lower energies with diffusion approach
- Monte Carlo simulation framework
- Models individual CRs or phase-space elements
- Includes many of the relevant interactions

CRPropa – A Community Driven Project

- Open source project with contributions from many institutes
 - Hamburg, Nijmegen, Madrid, Abu Dhabi, Aachen, Wuppertal, Bochum, L'Aquila, Bonn, Innsbruck, Paris, Oxford, Brussels, ...
- Mainly written in c++ with Python steering
- Works on simple laptops but scales well on computing clusters
- Easy to customize and extend

Basic Principle – The Candidate

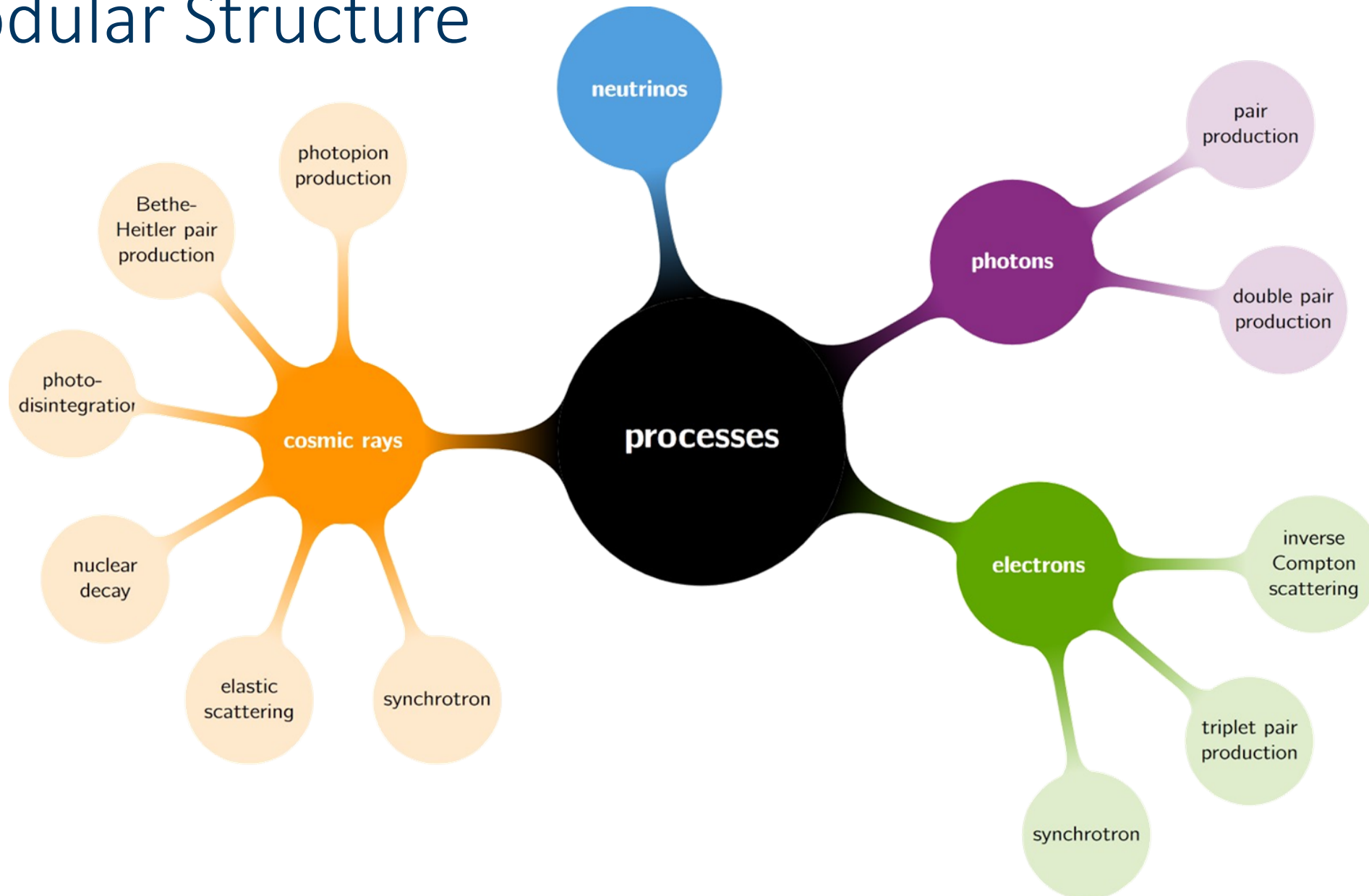


All modules must be designed stateless

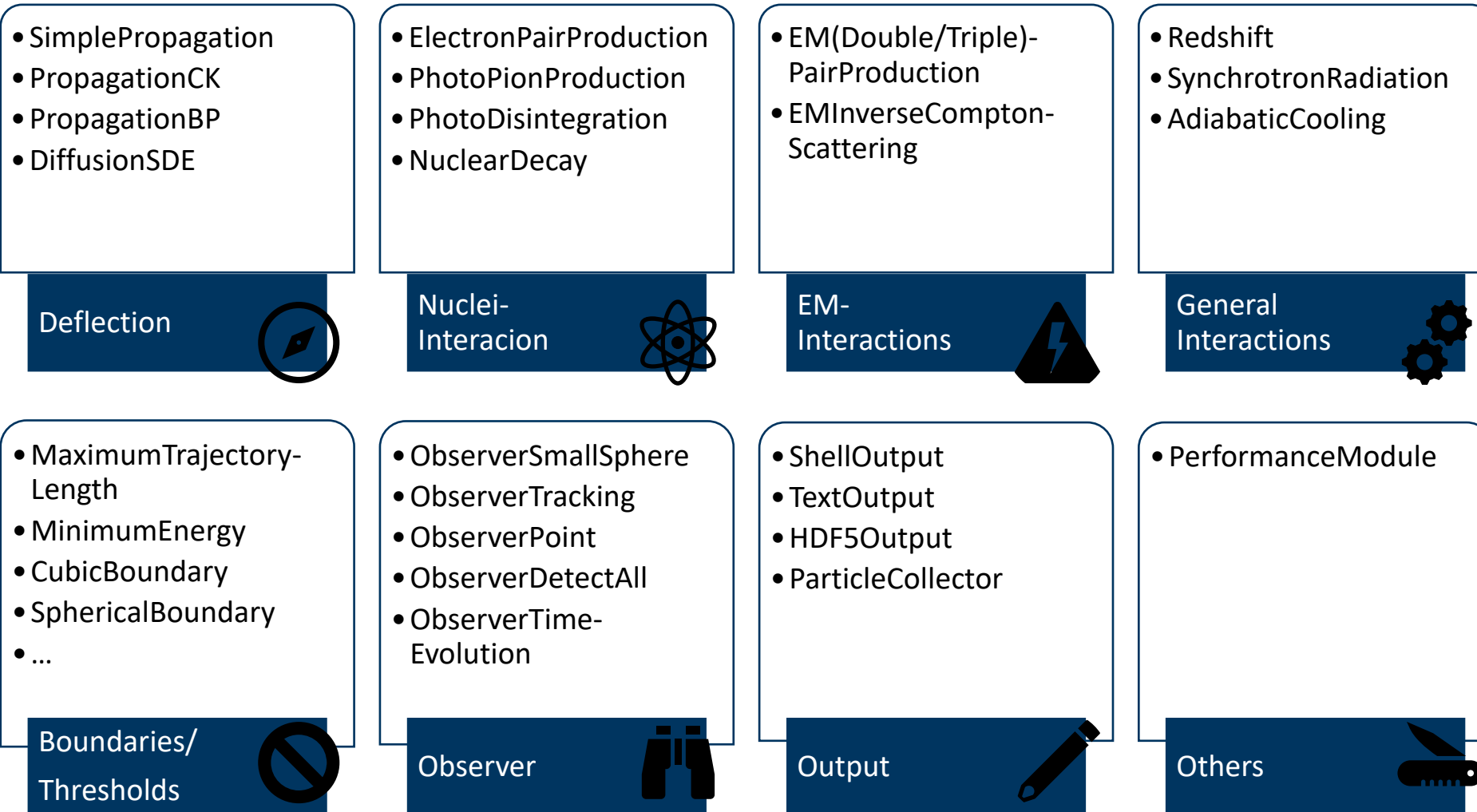


CRPropa can be parallelized without cross node communication!

Modular Structure

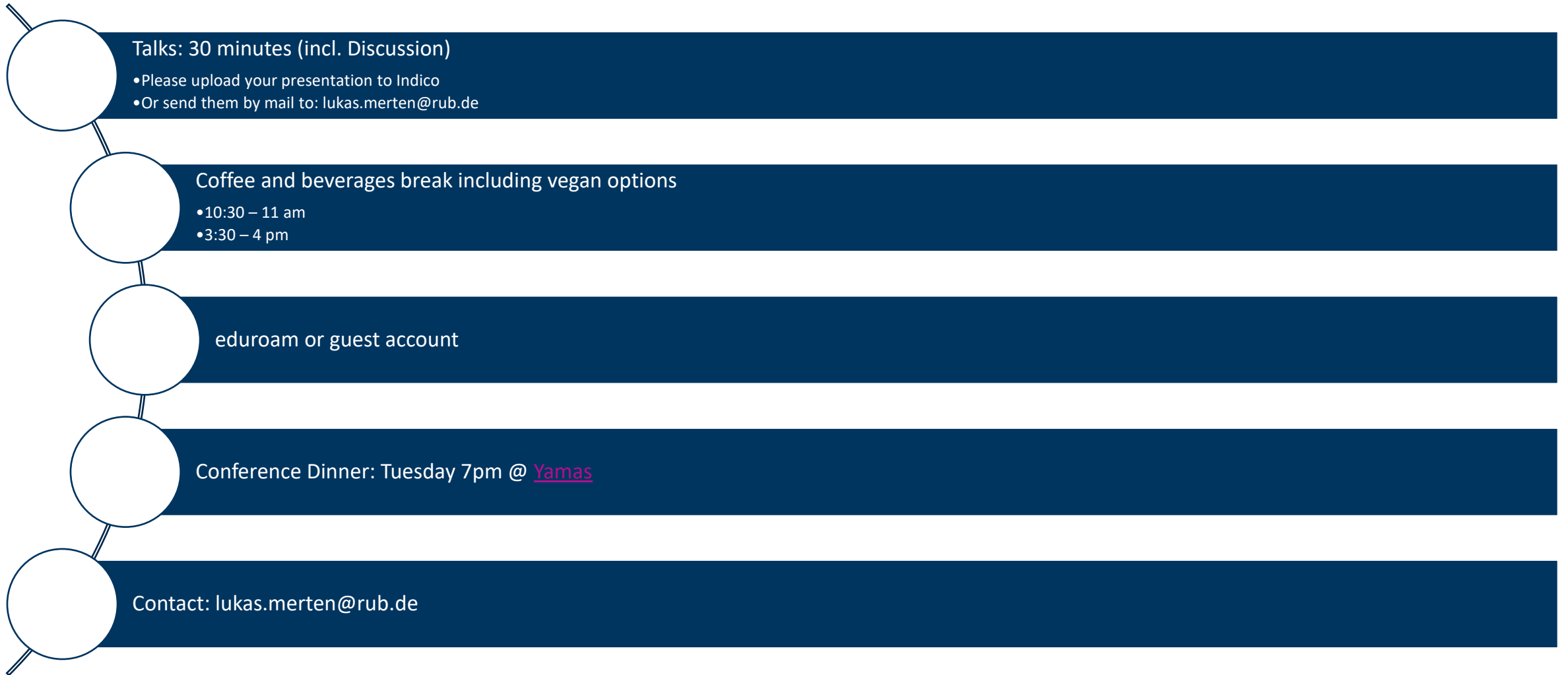


How to build a simulation?

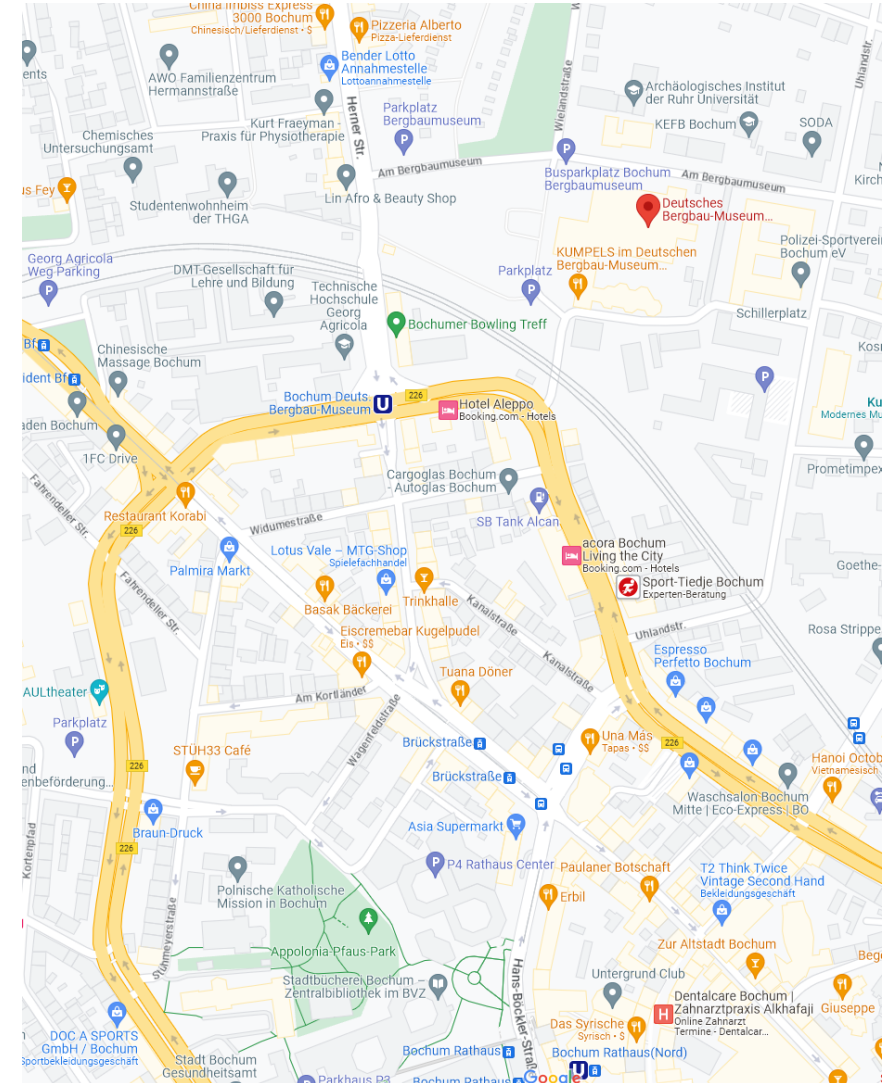
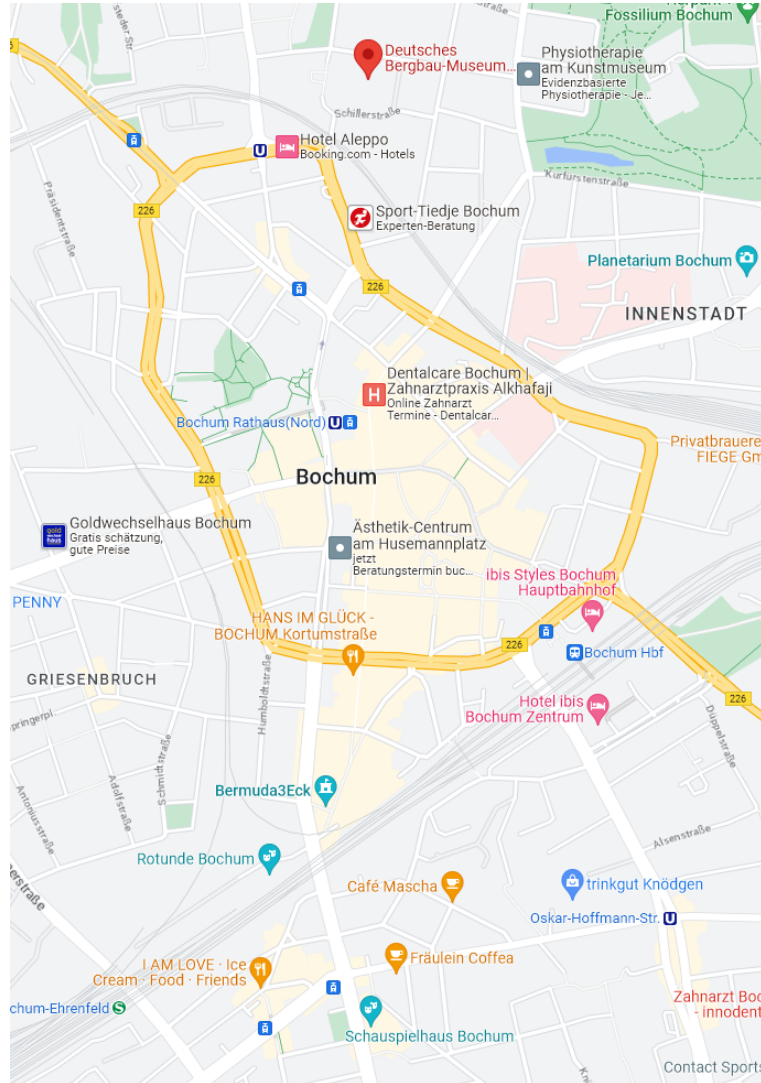


Organisation

Useful Information



Directions



Schedule

	Monday, September 25th		Tuesday, September 26th		Wednesday, September 27th	
Time	Speaker	Title	Speaker	Title	Speaker	Title
09:00-10:30	Lukas Merten	Introduction			Lukas Merten	Ensemble Averaged Transport with CRPropa
	Simone Rossoni	CRPropa simulations or anisotropy studies	Leander Schlegel	Time dependent particle injection to investigate the local source behaviour of flaring AGN	Horst Fichtner	t.b.a.
	Janning Meinert	Impact of a Λ CDM extension of UHECR propagation	Athithya Aravinthan	Synchrotron Radiation in CRPropa	Frederic Effenberger	Investigating Charged Particle Transport in Non-Gaussian Magnetic Turbulence Models
10:30-11:00	Coffee Break		Coffee Break		Coffee Break	
11:00-12:30	Frederik Krieger	Normalizing Flows for Parameter Estimation of UHECR Sources with Surface Detector Data of the Pierre Auger Observatory	Andrey Saveliev	On numerical Simulations of Intergalactic Electromagnetic Cascades with Lorentz Invariance Violations Using CRPropa	Pedro de la Torre Luque	The Galactic diffuse gamma-ray and neutrino emission at the PeV frontier
	Domenik Ehlert	Population studies of UHECR sources	Rafael Alves Batista	Coupling axion-like particles to photons during propagation: the ALPInist plug-in	Sophie Aerdker	Modeling Superdiffusive Particle Transport with CRPropa 3.2
	Pavlo Platko	Combined fit of the UHECRs with jetted AGN	Gaetano Di Marco	Characterisation of gamma-ray propagation in our Galaxy	Discussion	
12:30-14:00	Lunch Break		Lunch Break		Lunch Break	
14:00-15:30	Carmelo Evoli	Towards SimProp-Sirente: the new release of the SimProp Code (online)	Paolo Da Vela	Probing the intergalactic magnetic field through gamma-ray observations	Sebastian Hutschenreuter	The IMAGINE magnetic field library
	Discussion		Alexander Korochkin	Very-high-energy gamma ray propagation with CRPropa, Crbeam, and ELMAG	Discussion: The future of CRPropa	
15:30-16:00	Coffee Break		Coffee Break		Coffee Break	
16:00-17:30	Julien Dörner	Hadronic Interactions in CRPropa	Luis Enrique Espinosa	Influence of plasma instabilities on the propagation and spectrum of extragalactic electromagnetic cascades	Developer Meeting	
	Leonel Morejon	Simulating UHECR sources with CRPropa using updated interaction modules	Marcel Schroller	Local Source Physics with CRPropa - Extending Test Particle Simulations for Modelling Leptohadronic AGN Jets		
	Discussion		Discussion			
19:00			Dinner (YAMAS)			

Support

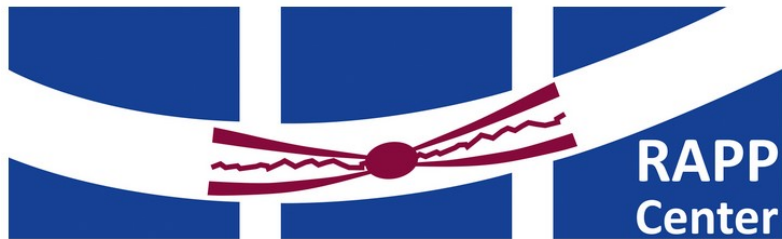
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Plasmas with Complex Interactions

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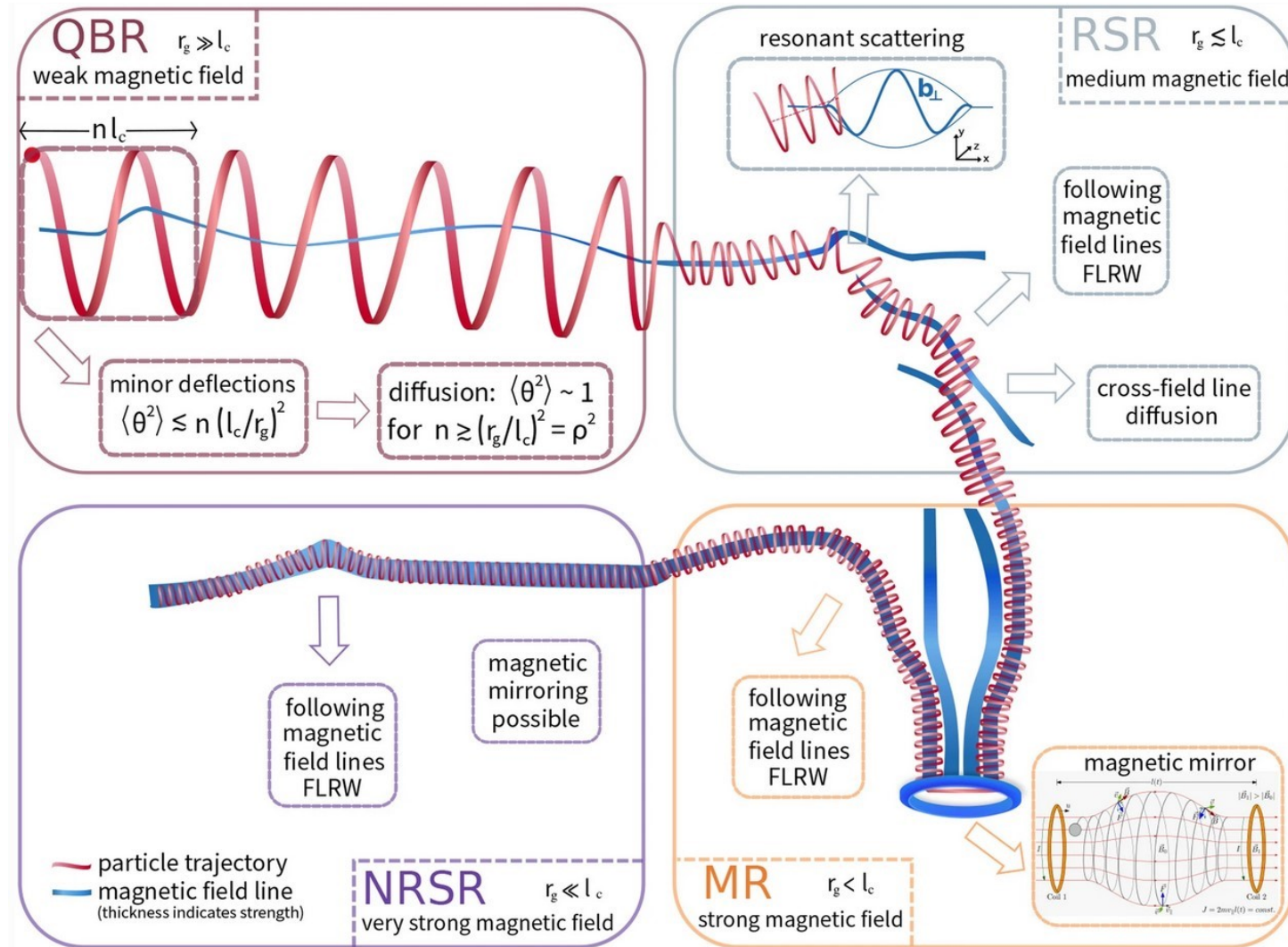
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Backup

Why do we need Modelling of CR Transport?

Magnetic Field Deflections



Magnetic Field Properties

- Can be structured on large scales
 - Direction, ordering
 - Strength
- Turbulent components are often observed
 - Power spectrum
 - Anisotropies
 - Intermittency

Random deflections

→ Smearing of sources

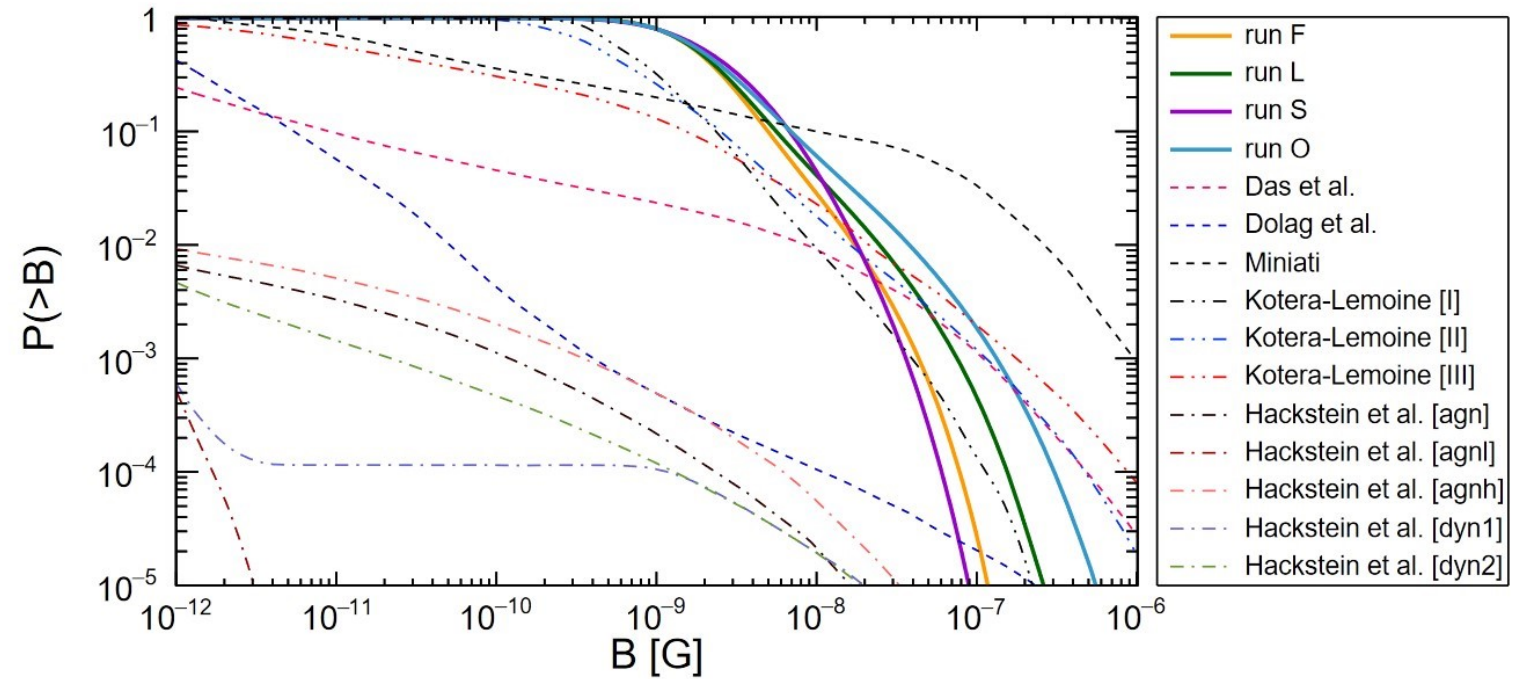
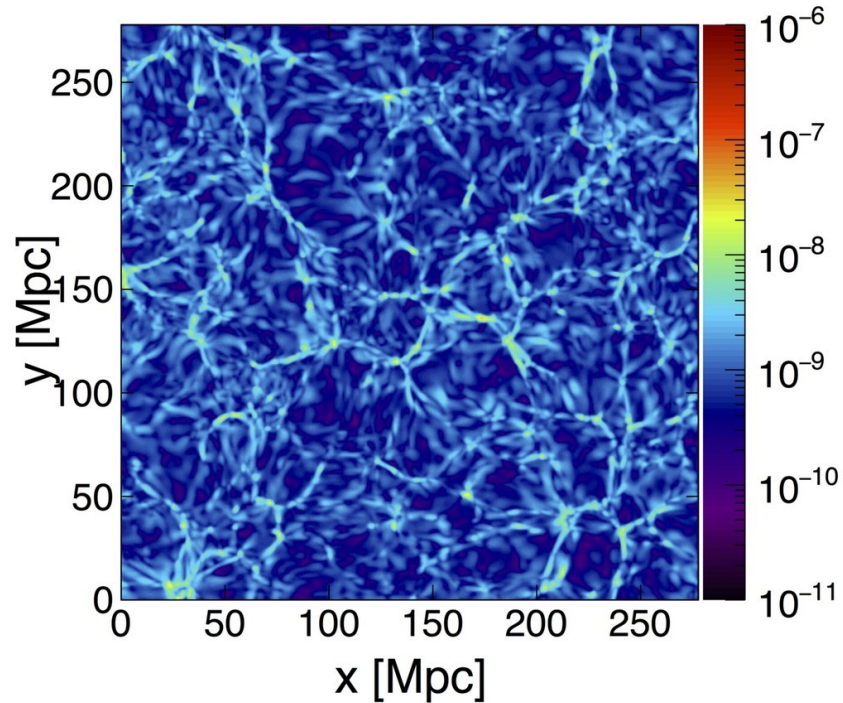
Coherent deflections

→ Shift of apparent source position

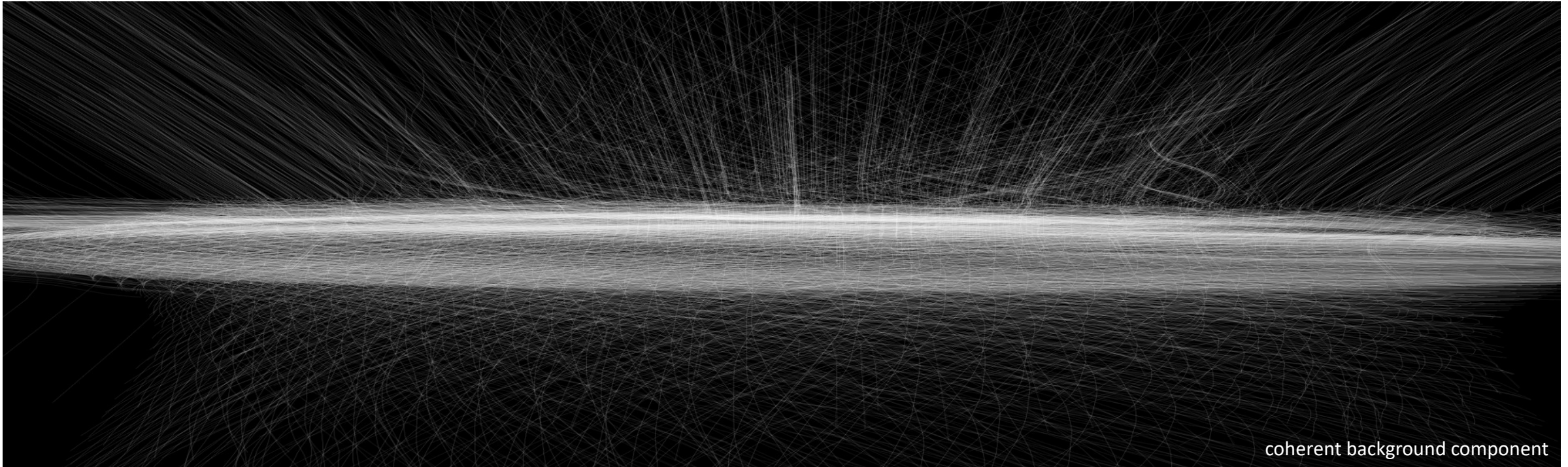
Combination

→ Magnetic horizon

Extra-Galactic Magnetic Fields



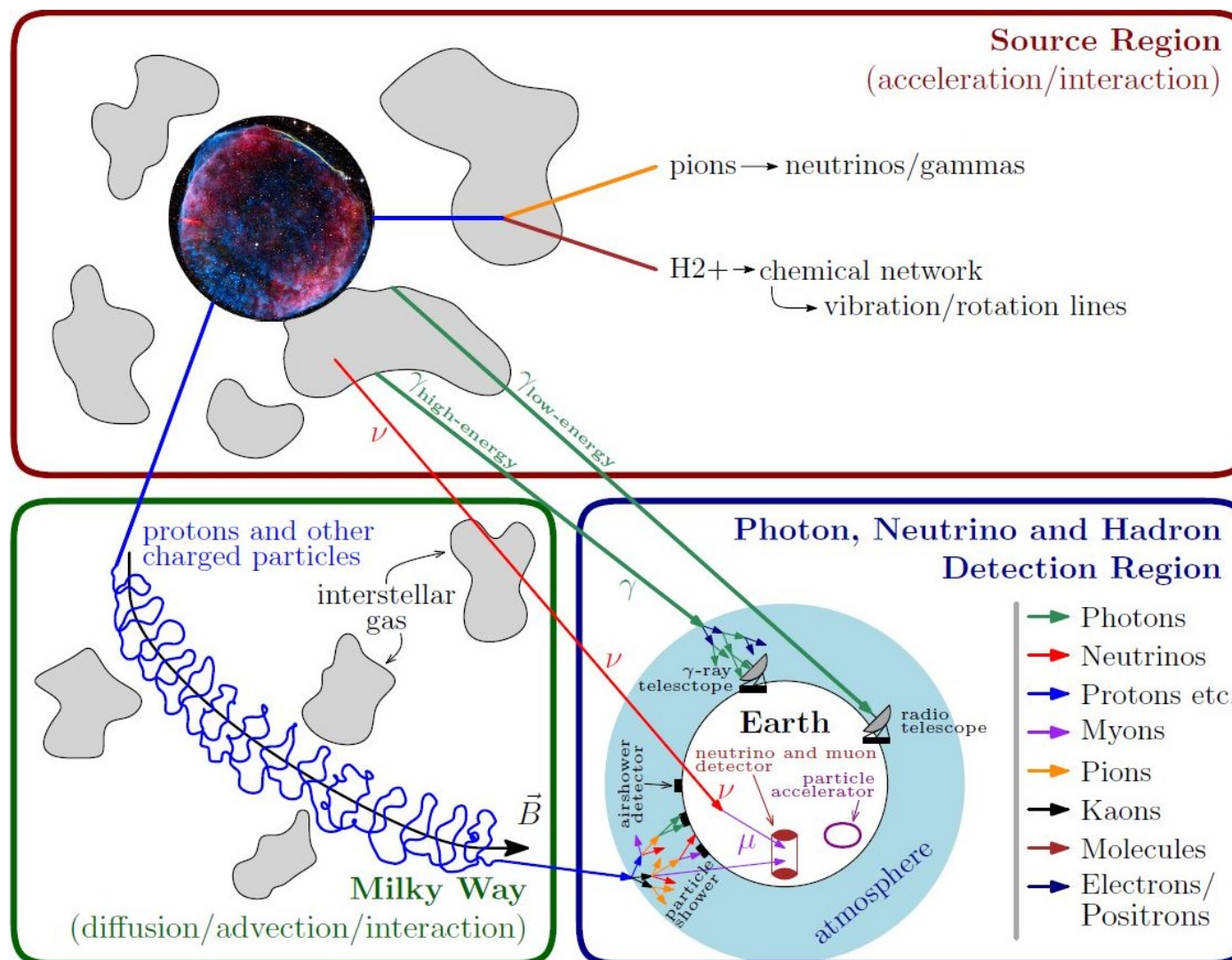
Galactic Magnetic Field



coherent background component

Why do we need Modelling of CR Interactions?

Interactions and Secondaries



Interactions and Losses

Cosmic Rays are never propagating in a pure vacuum

→ Energy losses and production of secondaries

Possible processes

Synchrotron Radiation, Ionization, Inverse Compton Scattering, (Double, Triple) Pair Production, Photo Meson Production, Photo Disintegration, Spallation, Nuclear Decay, ...

Possible Targets

- Photon Fields (CMB, EBL, Synchrotron photons, Corona, etc.)
- Hadronic Targets (HI/HII, H+, dust, He, etc.)
- Magnetic Fields

Secondaries

...everything that is produced during the propagation of CRs.

- Photons, Neutrinos, Electrons, Muons, lighter Cosmic Rays, etc.
- Can have complicated energy spectra
- Induce interactions themselves

Provide valuable otherwise inaccessible knowledge about CRs

- Source position ← neutrinos
- Energetics ← gamma rays
- Transport properties ← light cosmic rays (B/C ratio)
- Composition ← Muons